

BeppoSAX observation of the cluster Abell 970

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Abstract. We report new results on the cluster of galaxies Abell 970 obtained from X-ray observation with *BeppoSAX*. Our analysis of the *BeppoSAX* MECS and LECS data in the range [0.15–10] keV reveals a mean cluster gas temperature of $kT = 4.46^{+0.14}_{-0.15}$, a metallicity of $Z = 0.31^{+0.05}_{-0.04} Z_{\odot}$, and an interstellar hydrogen absorption column density of $N_{\text{H}} = 6.05^{+1.29}_{-0.97} \times 10^{20} \text{ cm}^{-2}$. Moreover, we obtained azimuthally averaged radial profiles of these quantities. Our results are consistent with the hypothesis that Abell 970 has been disturbed by a past merger or by the ongoing merger process of a substructure, that put the cluster out of equilibrium. This is also demonstrated by the offset between the gas and galaxy distributions. Combining the X-ray data with a recently published analysis of new galaxy radial velocities, we conclude that a subcluster 8 arcmin to the NW is falling into Abell 970 and will merge in a few Gigayears, thus disturbing Abell 970's newly acquired equilibrium. The high α -elements/iron ratio that we derive for this cluster supports the hypothesis of early intracluster medium enrichment by Type II supernovae.

Key words. galaxies: clusters: individual: Abell 970 – X-rays: galaxies: clusters

1. Introduction

Rich clusters of galaxies are the last structures to collapse in the Universe. When isolated and within its virial radius, the cluster contents should already be in quasi-stationary equilibrium. However, according to the hierarchical structure formation scenario, rich clusters are formed by the accretion of dark matter dominated haloes (White & Rees 1978 and, e.g., the N -body simulations of Evrard et al. 2002). These recurring merger events will affect the cluster, moving it away from an equilibrium state. Therefore, the study of the morphology and dynamics of clusters of galaxies can give valuable information on structure formation in the Universe.

The hot ($kT \approx 2$ –10 keV) and tenuous (central density, $n_0 \approx 10^{-3} \text{ cm}^{-3}$) X-ray emitting gas found in rich clusters of galaxies is an excellent tool to probe the cluster properties and past history. As it was already recognized almost three decades ago (Lea et al. 1973), the short relaxation time of the gas, and the fact that the main emission mechanism is the optically thin thermal bremsstrahlung, make it an excellent tracer of the cluster gravitational potential well.

Besides the temperature and density, the X-ray observation of the intracluster medium (ICM) can be used to measure the metal abundance (hereafter metallicity) of mainly iron, but also of some α -elements. Such observations can contribute to our understanding of gas enrichment in clusters, which depends on

the early evolution of galaxies in clusters and their interaction with the intra-cluster medium.

Here, we present an X-ray study of Abell 970 ($z = 0.0587$, richness class $R = 0$, B-M type III) following the work of Sodré et al. (2001, Paper I hereafter).

This cluster is in a rather dense environment, being part of the Sextans supercluster (number 88 in the catalogue of Einasto et al. 1997, and number 378 in the catalogue of Kalinkov et al. 1998, which also includes Abell 978, 979, and 993). It has moderate X-ray luminosity and it was observed with the *Einstein* Imaging Proportional Counter (e.g. Jones & Forman 1999); it was also detected in the *ROSAT* Bright Survey Catalogue (Schwope et al. 2000), although a pointed observation has never been carried out with *ROSAT*. In the study of White et al. (1997) using *Einstein* IPC data, it has been found that Abell 970 has at most a weak cooling flow, of about $20 M_{\odot} \text{ yr}^{-1}$ (see also Loken et al. 1999).

In the dynamical analysis made in Paper I, we found that the galaxies in Abell 970 have a mean velocity dispersion of 845 km s^{-1} (and up to $\sim 1000 \text{ km s}^{-1}$ in the centre). This analysis also indicated that this cluster must have substructure and is out of dynamical equilibrium. This conclusion is also supported by an offset in the peaks of the surface density distribution and the X-ray emission obtained from an *Einstein* IPC image. Further support came from a large-scale velocity gradient in the cluster. We also found a discrepancy between the dynamical masses inferred from the virial theorem and those inferred with the X-ray emission, which is expected when the galaxies and the gas inside the cluster are not in virial (or

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hydrostatic) equilibrium. Here we present new results on the distribution of temperature, density, and abundances of the X-ray emitting gas, as well as its α -elements/iron ratio.

This paper is organized as follows. We describe in Sect. 2 the data obtained with *BeppoSAX*. In Sect. 3 we present the X-ray fluxes and luminosities and determine and analyse the temperature and metallicity radial profiles. The cluster gas and dynamical masses are estimated in Sect. 4. In Sect. 5 we discuss the possible merging scenario of Abell 970, and in Sect. 6 we give the individual metal abundances and discuss the consequences for the ICM enrichment modelling.

2. The data

Abell 970 was observed in July 2000 by the *BeppoSAX* satellite (Boella et al. 1997a) with two of the narrow field instruments: the LECS providing sensitivity in the [0.1–4.0] keV band (Parmar et al. 1997), and the MECS 2 and 3, which are sensitive in the [1.6–10.5] keV range (Boella et al. 1997b). The net exposure times were 97 289 and 33 619 s for the MECS and LECS, respectively.

We restricted the scientific analysis to the 8 arcmin ($\sim 760 h_{50}^{-1}$ kpc at the distance of Abell 970) circle centered at the aim point of the LECS and MECS instruments in order to prevent severe vignetting effects (Cusumano & Mineo 1998). Moreover, the MECS entrance window is supported by a beryllium “strongback” that is opaque to low energy photons ($E \lesssim 5$ keV) and is not taken into account in the effective area files (Fiore, private communication). The strongback structure affects detector positions between 7.9' to 13.3' (Cusumano & Mineo 1998). Since we are analysing data inside the MECS window structure, we opted for combining both MECS 2 and 3 event files in order to improve the signal-to-noise ratio, instead of analysing them separately as De Grandi & Molendi (2002), which would be essential for an analysis extending beyond $\sim 8'$.

Figure 1 shows the X-ray isocontours superposed on a “Second Epoch Survey” DSS optical image¹. The event file coordinates were corrected for the systematic error reported by the *BeppoSAX* team². After correction, we verified a very good matching of both the MECS and LECS images to the Einstein IPC image (see Paper I), including the small X-ray peak at the SE, which, as pointed out in Paper I, is associated to a pair of interacting cluster galaxies (numbers 37 and 38 of Paper I catalogue: $\alpha = 10^{\text{h}}17^{\text{m}}37^{\text{s}}$, $\delta = -10^{\circ}46'02''$, J2000). The distance between the optical position of the pair and the associated X-ray peak is less than 15''.

A comparison of Fig. 1 with the *Einstein* IPC image (Fig. 4 of Paper I) demonstrates that most of the features seen here were already present in the IPC image. The peculiar offset between the X-ray and the galaxy distribution maxima noticed before – about 0.8 arcmin – is also evident. As discussed in Paper I, this offset is a consequence of the displacement of the

X-ray isophotes in the direction of a concentration of galaxies to the northwest ($\alpha \sim 10^{\text{h}}16^{\text{m}}48^{\text{s}}$, $\delta \sim -10^{\circ}38'$) which is centered on the brightest cluster galaxy. This galaxy group however has no associated X-ray excess (see also Fig. 7, below).

Given the moderate spatial resolution of the MECS and LECS detectors ($FWHM$ 1.4 arcmin at 6 keV and 3.6 arcmin at 2 keV, respectively), we have defined sub-regions in the field of view in order to perform separate spectral analysis. The X-ray data in these regions were extracted with *xSELECT* 2, from the *HEASOFT* package³.

We have defined 6 concentric annuli, each 1'28'' wide (11 pixels), centered at the X-ray maximum ($\alpha = 10^{\text{h}}17^{\text{m}}21^{\text{s}}$, $\delta = -10^{\circ}41'32''$). These regions are employed to derive the radial profiles of various physical quantities. Notice that the last annulus extends over the detector region affected by the MECS “strongback”.

We have also defined 4 regions corresponding to the NW, NE, SW, and SE quadrants (excluding the central region, see Sect. 5 for details). This subdivision in sectors is very much like the one used by De Grandi & Molendi (1999). These regions, which are depicted in Fig. 7 below, were chosen based on the galaxy distribution and their velocity distribution (cf. Paper I). Finally, we have defined a circular region of 2 arcmin (15 pixels) radius centered at the peak of the SE X-ray excess ($\alpha = 10^{\text{h}}17^{\text{m}}37^{\text{s}}$, $\delta = -10^{\circ}46'02''$, see Fig. 1).

The X-ray spectrum of each region was modelled as being produced by a single temperature plasma and we employed the *MEKAL* (Kaastra & Mewe 1993; Liedahl et al. 1995) model. For the whole cluster data, we also used the *VMEKAL* model which allows the individual abundances for the elements to be fitted independently. However, the abundance of some metals cannot be constrained by the spectral fitting, therefore we have only used the following elements: Si, S, Ar, Ca, Fe, and Ni. The others metals were fixed to 0.3 Z_{\odot} .

The spectral fits were done using *xSPEC* v11.0. Following the *BeppoSAX* “Cookbook” (Fiore et al. 1999), we have selected the data in the energy range [0.12–4.0] keV for the LECS, and [1.65–10.5] keV for the MECS. The photoelectric absorption – mainly due to neutral hydrogen – was computed using the cross-sections given by Balucinska-Church & McCammon (1992), available in *xSPEC*. Figure 2 shows four examples of spectral fits.

For the last annulus, we have also done the spectral fits using only MECS data in the interval [5.0–10.5] keV, in order to check the masking effect of the “strongback”. We verified that the difference between the best fit values obtained either in the full energy range [1.65–10.5] keV or in the [5.0–10.5] keV range is well within the error bars (cf. below, Sect. 3.2).

Background spectra were constructed from the event files (produced by the *BeppoSAX* Science Data Centre) of the standard background data, which represent an assemblage of several empty fields devoid of any noticeable source. For the MECS, we have used the combined background for detectors 2 and 3 available from the *BeppoSAX* web site (under

¹ The “Second Epoch Survey” of the southern sky was made by the Anglo-Australian Observatory (AAO) with the UK Schmidt Telescope. Plates from this survey have been digitized and compressed by the STScI UK Schmidt Telescope and the Digitized Sky Survey.

² cf. www.asdc.asi.it/bebbosax/coord_correction.html

³ A service of the Laboratory for High Energy Astrophysics (LHEA) at NASA/GSFC and the High Energy Astrophysics Division of the Smithsonian Astrophysical Observatory (SAO).

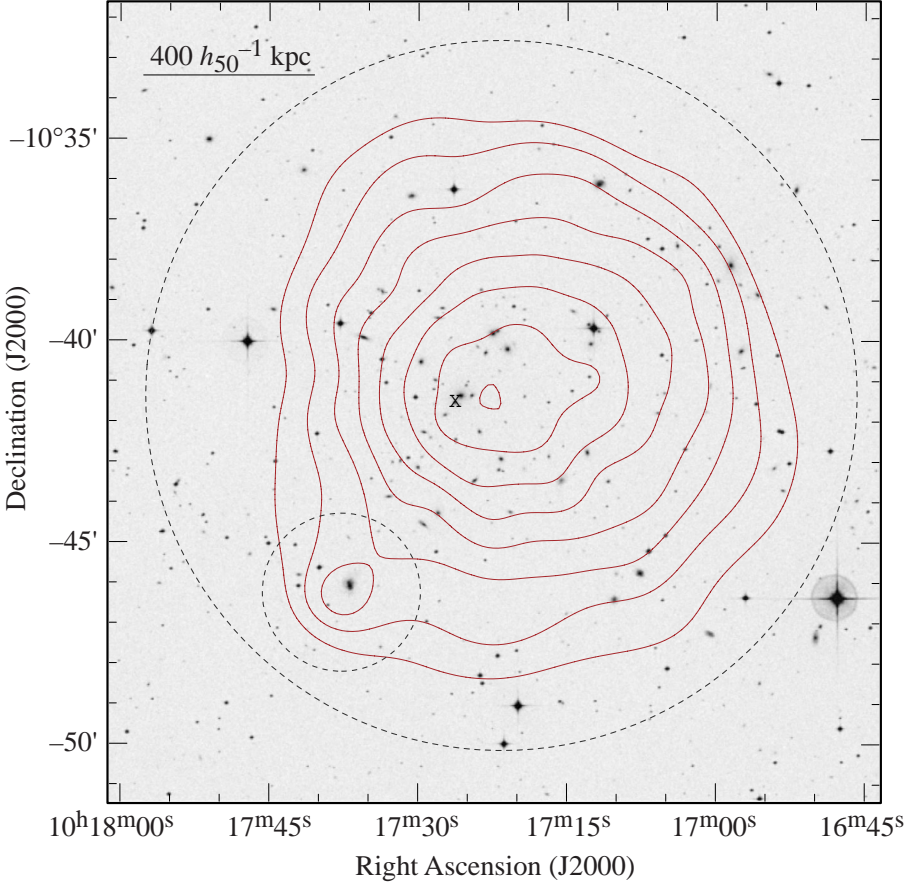


Fig. 1. X-ray isocontours obtained with *BeppoSAX* MECS 2 and 3 superposed on a red DSS image (shown in grey scale). The larger dashed circle indicates the area used in our analysis; the smaller dashed circle indicates the excess X-ray emission tentatively associated with an active galactic nucleus (see text). The contours are spaced logarithmically. Notice the offset between the X-ray maximum and the centre of the galaxy distribution ($\alpha = 10^{\text{h}}17^{\text{m}}26.4^{\text{s}}$, $\delta = -10^{\circ}41'27''$, J2000, marked with an “X”, see also Fig. 7, below).

/pub/sax/cal/bgd/98_11). For the LECS, we have used the background event file described by Parmar et al. (1999) and available at /pub/sax/cal/bgd/99_12. The background spectra were extracted at the same regions (in detector coordinates) defined above and used in the spectral fits.

3. Results

3.1. Flux and luminosity

Table 1 summarizes the fluxes and luminosities obtained for the whole usable field and the central region. The bolometric X-ray luminosity implies $T_{\text{X}} = 5.0 \pm 0.5$ keV based on the $L_{\text{X}}-T_{\text{X}}$ relation for clusters obtained by Xue & Wu (2000).

As suggested in Paper I, the SE X-ray emission excess visible in Fig. 1 may be due to an active nucleus excited by the merger of the pair of interacting galaxies. In order to verify this hypothesis, we have extracted a spectrum in a 2 arcmin circle comprising this feature and tested different spectral models (see Table 2).

We have tested a power-law (pow) and a thermal plasma (MEKAL) model. A composite model, power-law plus *mekal* (i.e., a non-thermal source superposed on the cluster thermal emission) could not be fitted unambiguously – fits with different initial parameters ended with different minimum χ^2 in parameter space. Therefore, we considered two opposite scenarios: I) the X-ray excess comes from a non-thermal source and the emission inside a circle of 2 arcmin will be dominated by

Table 1. Non-absorbed fluxes and luminosities in different energy bands. The energy band limits are given in keV.

Region	flux [0.1–2.4]	flux [0.5–4.5]	L_{X} [0.1–2.4]	L_{X} [0.5–4.5]	L_{X} bolom.
All field	27.8	30.0	42.2	46.3	83.1
Centre	5.13	5.69	7.78	8.75	16.3

Note: flux units are 10^{-12} erg cm^{-2} s^{-1} .
Luminosity units are 10^{43} h_{50}^{-2} erg s^{-1} .

Table 2. Spectral fits of the south-east X-ray excess. N_{H} is the hydrogen column density in units of 10^{20} cm^{-3} , kT is the temperature in keV, Z is the metallicity, α is the spectral index and “dof” is the degrees of freedom. Errors are at the 1σ confidence level.

Model	N_{H}	kT	Z/Z_{\odot}	α	χ^2/dof
POW	< 4.1	—	—	1.5 ± 0.1	80.7/89
MEKAL	< 2.5	15.2 ± 7.5	0.35^{\ddagger}	—	80.0/89
MEKAL	6.05^{\ddagger}	14.2 ± 6.6	0.35^{\ddagger}	—	81.1/90
MEKAL	< 50	4.5^{\ddagger}	0.55 ± 0.29	—	94.1/89
MEKAL	6.05^{\ddagger}	4.5^{\ddagger}	0.58 ± 0.33	—	95.3/90

\ddagger Value fixed.

a power-law spectrum; II) the merging galaxies have no AGN and the emission is mostly of thermal origin.

We found that the best models, both statistically equivalent, are the power-law and the thermal plasma endowed with

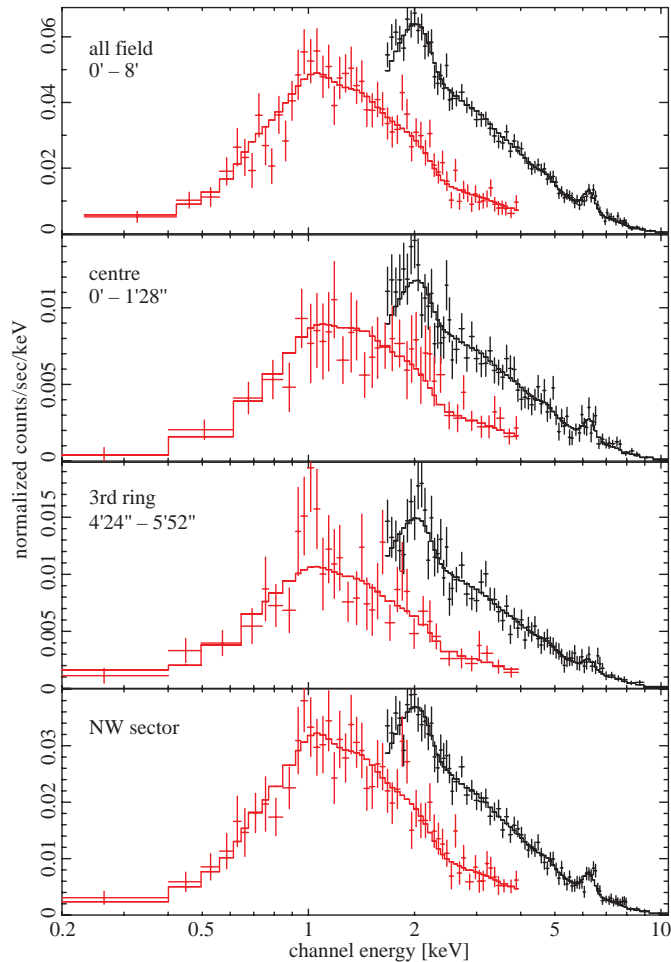


Fig. 2. Four examples of the spectral fit. The left spectrum in each panel is from the LECS, the right one from the combined MECS 2 and 3. The lines passing through the data points correspond to the best fit absorbed MEKAL model.

the unreasonably high temperature of 14 keV or more. Fixing the plasma temperature at 4.5 keV (i.e., about the cluster mean temperature, see below), whether or not keeping fixed metallicities and N_{H} , produces poorer fits. We are thus led to suggest that the X-ray source is of non-thermal origin, probably due to the nuclear activity of one or both of the components of the interacting pair of galaxies optically associated with it. This conclusion should however be regarded with caution given the poor quality of the spectrum in this region, with net count rates of 5.1×10^{-3} count/s for the MECS and 2.3×10^{-3} count/s for the LECS.

3.2. Radial profiles

We have computed the radial profiles for several astrophysical quantities using the circular annuli defined in Sect. 2, with the combined MECS and LECS data. Since the SE X-ray excess may be of non-thermal origin, we have done the spectral fits excluding and including this region. The results are virtually the same, because the contribution of the SE excess to the X-ray counts inside the annuli is not significant. The X-ray SE excess was therefore excluded from the data prior to the spectral

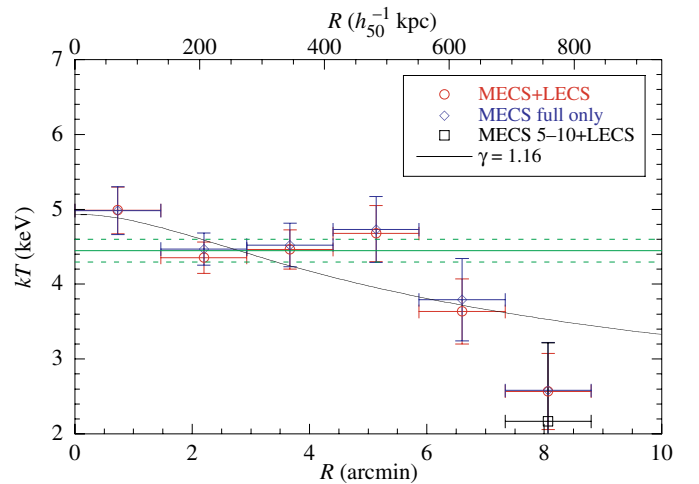


Fig. 3. Temperature profile using either both LECS and MECS or only MECS data (which is more appropriate for cluster temperatures $T \gtrsim 3$ keV). For the last annulus there is an extra data point (square) where the spectral fit was done using only the MECS [5.0–10.5] keV band. The error bars are at the 1σ confidence level. The horizontal lines show the mean temperature (full line) and the 1σ error bar (dashed) measured using all data within 8 arcmin ($kT = 4.46^{+0.14}_{-0.15}$). The curve (full line) shows the best fit of a polytropic model (see text).

fits presented here. Figure 3 shows the resulting temperature profile⁴.

The temperature profile was also determined using only MECS data and fixing $N_{\text{H}} = 5.42 \times 10^{20}$ cm⁻² (the galactic value, Dickey & Lockman 1990). This was done for two reasons: first, to check the robustness of the fits, whether or not there was any spurious effect introduced by the use of two different detectors.

The second reason is that when both the temperature and hydrogen column density are taken as free parameters, and depending on the spectral range being adjusted, the spectral fitting procedure may produce spurious correlations between best-fit values of kT and N_{H} (see, e.g., Pislari et al. 1997).

As it can be seen from Fig. 3, the resulting temperature profiles, that is, for MECS data only and N_{H} fixed at the galactic value and for the combined MECS and LECS data and N_{H} free, are almost indistinguishable. This demonstrates the internal consistency of the modelled temperatures and a spurious correlation between fitted kT and N_{H} values is not present.

An additional fit for the outermost annulus was done using only the [5.0–10.5] keV MECS data (but still using the full LECS). This was done to verify whether or not the MECS “strongback” has an important masking effect for low energy photons. It turns out that the effect on the temperature determination is within the error bars (which increases compared to the fit using the full MECS range).

A simple parametric form of the ICM temperature can be obtained using the polytropic equation of state. A polytropic

⁴ Actually, it is the *emission weighted temperature* that is obtained in a spectral fit. However, for physical realistic gas models, the difference between the actual gas (electron) temperature and the measured emission weighted temperature should be negligible and well within the observational uncertainties (Komatsu & Seljak 2001).

model for the ICM is motivated by the results of hydrodynamical simulations (e.g., Suto et al. 1998), theoretical models (e.g., Cavaliere et al. 1999; Dos Santos & Doré 2002), and observed temperature gradients (e.g., Markevitch et al. 1999; Finoguenov et al. 2001a, but see also Irwin et al. 1999; Irwin & Bregman 2000; De Grandi & Molendi 2002 for a different observational result – see discussion below).

The polytropic gas is defined by:

$$P \propto \rho^\gamma, \quad (1)$$

where P and ρ are the pressure and density of the gas, respectively, and γ is the gas polytropic index. For an ideal gas, $P = nkT$ (n is the number density of particles), therefore the temperature and density are related by:

$$kT \propto n^{\gamma-1}. \quad (2)$$

For the gas density profile, we use the well known β -model (Cavaliere & Fusco-Femiano 1976):

$$n(r) = n_0 \left[1 + \left(\frac{r}{r_c} \right)^2 \right]^{-3\beta/2}, \quad (3)$$

where n_0 , r_c and β are the central density, the core radius, and the shape parameter. We adopt the values determined in Paper I: $\beta = 0.66 \pm 0.44$ and $r_c = (260 \pm 20)h_{50}^{-1}$ kpc or 2.8 ± 0.2 arcmin. For our considerations we have used the *Einstein* IPC results because this instrument provides a higher spatial resolution (about 25 arcsec) compared to the *BeppoSAX* MECS (about 1.5 arcmin). The temperature profile is then written as:

$$T(r) = T_0 \left[1 + \left(\frac{r}{r_c} \right)^2 \right]^{-3\beta(\gamma-1)/2}, \quad (4)$$

where T_0 is the central temperature.

A least-square fit of Eq. (4) yields $\gamma = 1.15 \pm 0.05$ and $T_0 = 4.93 \pm 0.23$ with $\chi^2 = 9.3$ for 4 degrees of freedom. This value agrees well with those predicted by Komatsu & Seljak (2001) and found by Finoguenov et al. (2001b). The fit, shown in Fig. 3, is poor because of the very steep temperature gradient beyond 5 arcmin. In their study of a sample of 21 nearby clusters observed with *BeppoSAX*, De Grandi & Molendi (2002) find that the temperature profiles of clusters are characterized by isothermal cores extending to about one-fifth of the cluster virial radius, followed by steep temperature gradients which are steepest for the non-cooling flow clusters. These authors also find poor fits when trying to adjust a polytropic law to the overall temperature profiles. This situation is much similar to that observed in Abell 970 which, when compared to their sample of non-cooling flow clusters, presents a steeper gradient. In Sect. 4.2 we use both the isothermal and the polytropic model to estimate the total cluster mass.

We may check whether the last three radial points still correspond to a stable state or whether the temperature gradient is too steep for stability. If we take into account only the three external points, i.e., $5 < R < 8$ arcmin, the polytropic index increases to $\gamma = 1.7 \pm 0.2$, which is slightly above the ideal gas value, $5/3$, suggesting that the outer region may be barely in stable equilibrium.

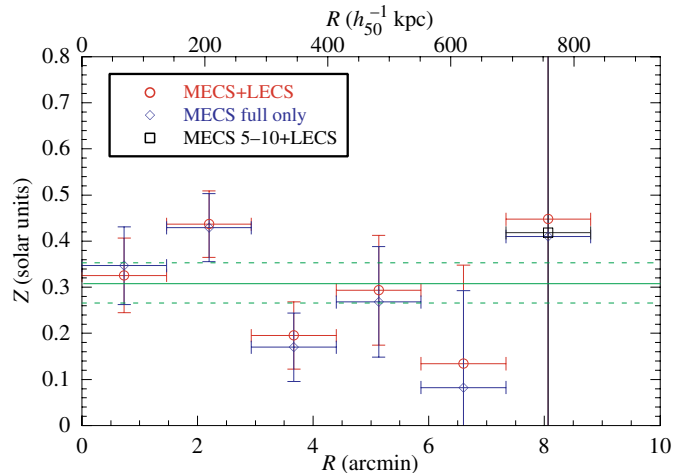


Fig. 4. Metallicity profile using either both LECS and MECS or only MECS data (the metallicity is determined mainly by the iron complex at ~ 6.8 keV). The square symbol has the same meaning as in the previous figure. The error bars are 1σ confidence level. The horizontal lines show the mean metallicity measured using all data inside 8 arcmin ($Z = 0.31_{-0.04}^{+0.05} Z_{\odot}$).

The *MEKAL* plasma model spectral fit also gives the ICM metallicity. Figure 4 shows the metal abundance radial profile. In order to check the fit robustness, we also did the profile determination using only the MECS data, besides the fits using both detectors. Both profiles are very similar, well within the error bars.

As before, we made an additional fit for the last annulus using only photons with $E \geq 5$ keV. Although the metallicity value obtained is close to the spectral fit with the full MECS range, the error bars are so big that, in fact, the metallicity should be considered undetermined at this last annulus.

The metallicity profile shown in Fig. 4 has no apparent gradient, being consistent with the mean value found for the whole cluster (which is a typical value for a cluster of such richness and temperature). Contrary to some other clusters, e.g., Abell 496 (Dupke & White 2000 using *ASCA* data) or Abell 85 (Lima Neto et al. 2001, using *BeppoSAX* data), this cluster does not show an increase of metal abundance in the centre (defined by the X-ray peak emission). Irwin & Bregman (2001), also using *BeppoSAX* data, find that clusters generally show a negative metallicity gradient, but this gradient is usually smaller or even disappears in non-cooling flow clusters.

Figure 5 shows the radial profile of hydrogen column density. For comparison we also show the galactic value at the position of Abell 970. Here we observe a gradient, with significantly higher hydrogen column density in the inner 2 arcmin. This excess is significant at a 2σ level, ignoring the uncertainty of the galactic N_{H} measure.

This level of significance may indicate that we are detecting an X-ray absorption excess compared with the expected absorption predicted by the observed galactic HI column density. The origin of such an excess is still controversial (Allen 2000, and references therein), the most promising hypotheses being very cool molecular gas and/or dust grains. Moreover,

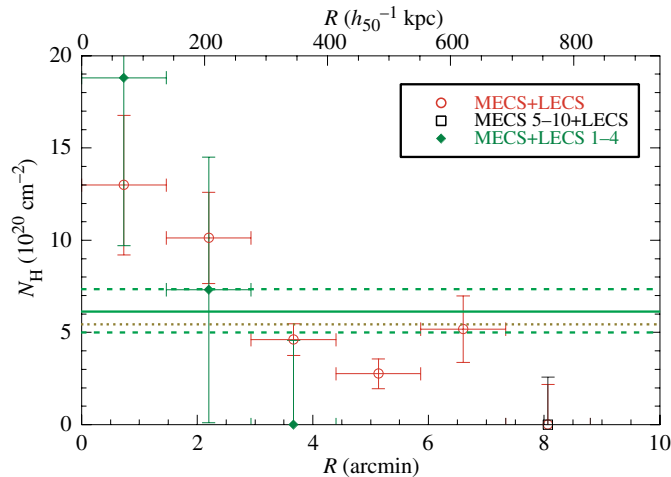


Fig. 5. Hydrogen column density profile using LECS and MECS data. The error bars are 1σ confidence level. Filled diamonds correspond to fits with LECS counts restricted to the range [1–4] keV (but all MECS). The open square corresponds to the last annulus fit in which MECS counts were restricted to the range [5–10.5] keV (see text). The horizontal full line shows the mean hydrogen column density measured using all data within 8 arcmin; the horizontal dashed lines correspond to 1σ error bars ($N_{\text{H}} = 6.05^{+1.29}_{-0.97} \times 10^{20} \text{ cm}^{-2}$). The horizontal dotted line indicates the mean galactic value, $N_{\text{H}} = 5.42 \times 10^{20} \text{ cm}^{-2}$.

the amount of this cool material depends on the details of the spectral modelling (McCarthy et al. 2002).

Figure 5 shows that N_{H} decreases outwards, reaching the galactic value. However, a caveat must be introduced here: (I) there may be variations of the galactic N_{H} that are unresolved by the galactic survey at the cluster scale; (II) the N_{H} determination relies almost exclusively on the LECS data (photon energy $\lesssim 1.5$ keV). The error bars in this case may be underestimated given the low count rate, specially at the last bin.

Moreover, there are a number of issues associated with lowest energy photons detected by the LECS: the spatial resolution is poor ($\approx 5.14/\sqrt{E_{\text{keV}}}$ arcmin) and there is a disagreement between the measured (on-ground) and theoretical effective area for $E \lesssim 0.3$ keV (see Parmar et al. 1997). Selecting only LECS counts with $E \geq 1.0$ keV (the spatial resolution is ~ 5 arcmin at $E = 1$ keV) can minimize these effects. Figure 5 shows the additional N_{H} determination with such a selection, for the centre and the first two rings. The error bars are significantly larger, but the central excess in modelled neutral hydrogen column density is still present at 2σ significance.

4. Cluster mass

4.1. Gas mass

In order to determine the cluster gas mass we assume that the gas distribution can be described by a spherically symmetric *beta*-model profile with values of core radius, r_c , and shape parameter, β as determined in Paper I. With the good quality spectrum we now dispose, it is possible to estimate the central electronic density, n_0 , as follows.

We integrate the bremsstrahlung emissivity along the line-of-sight, in the energy bands given in Table 1 and obtain n_0 .

Taking into account the error bars in r_c and β , we obtain $n_0 = (3.7 \pm 0.6)10^{-3} \text{ cm}^{-3}$. Notice that the shape of the temperature profile has only a minor effect (less than 5%) in the determination of n_0 .

The integrated gas mass profile (or growth curve) has been obtained by integration of Eq. (3) and is shown below in Fig. 6.

4.2. Dynamical mass

With the density and temperature profiles, we can estimate the total cluster mass. Supposing the ICM is in hydrostatic equilibrium, the total “X-ray” dynamical mass (i.e., the dynamical mass determined from an X-ray observation, not to be confused with the mass of the X-ray emitting gas) inside r is given by:

$$M(r) = -\frac{kT}{G\mu m_p} r \left(\frac{d \ln \rho(r)}{d \ln r} + \frac{d \ln T(r)}{d \ln r} \right), \quad (5)$$

where μm_p is the mean mass per particle ($\mu = 0.59$ for a fully ionized primordial gas) and m_p is the proton mass. We assume that matter is distributed spherically.

Adopting the β -model to describe the intracluster gas density and the polytropic equation of state, the dynamical mass results in:

$$M(r) = \frac{3kT_0\beta\gamma r_c}{G\mu m_p} \left(\frac{r}{r_c} \right)^3 \left(1 + \left[\frac{r}{r_c} \right]^2 \right)^{-1-\frac{3}{2}(\gamma-1)\beta}. \quad (6)$$

When $\gamma = 1$ we have the usual expression for the isothermal sphere. Figure 6 shows the mass profile computed both with an isothermal temperature and polytropic model.

The dynamical masses estimated here are about 15% to 25% higher than those given in Paper I where a 3.3 keV isothermal model was assumed.

We can now compute the radius corresponding to the density ratio $\bar{\rho}(r_0)/\rho_c(z)$ (see Appendix). For $\delta = 200$, we have the usual r_{200} that is often used as an approximation to the virial radius (e.g., Lacey & Cole 1993). Assuming either an isothermal or polytropic profile, we obtain $r_{200} = 2.3 h_{50}^{-1} \text{ Mpc}$ and $r_{200} = 1.9 h_{50}^{-1} \text{ Mpc}$, respectively. Thus, R_{image} , the maximum radius analysed, is $0.33\text{--}0.4 r_{200}$, depending on the adopted temperature profile.

From the dynamical and gas mass profiles, we can derive the baryon fraction profile, $f(r) \equiv M_{\text{baryon}}/M_{\text{total}}$, assuming that the bulk of the baryons are in the cluster ICM (i.e., $M_{\text{baryon}} \approx M_{\text{gas}}$). The neglect of the galactic contribution (about ~ 5 to 10% of the total mass, see Paper I), underestimates f by just a few percent.

The bottom panel of Fig. 6 gives the estimated baryon fraction profile. For reference, we show r_{2500} and r_{5000} , the radii within the mean cluster density is 2500 and 5000 times the value of the critical density of the Universe.

The gas mass fraction obtained here agrees with the typical values obtained by White & Fabian (1995) (at $R = 1 \text{ Mpc}$) and by Allen et al. (2002) (at $R = r_{2500}$). However, the shape of the $f(r)$ is much flatter than the profiles computed by Allen et al. A possible explanation might be the adopted model for the mass: while we assume a β -model for the gas profile, Allen

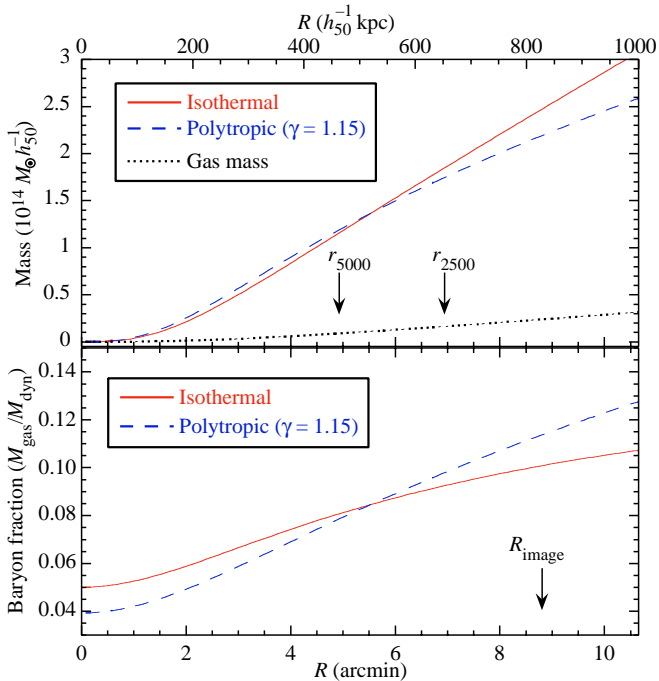


Fig. 6. Top: X-ray dynamical mass profile using an isothermal model, with the mean cluster temperature (solid line), and polytropic model, with central temperature $T_0 = 5.0$ keV and $\gamma = 1.15$ (dashed curve). The gas mass is also shown (dotted curve). Bottom: Baryon fraction given by the dynamical and gas mass ratio. It is shown r_{2500} and r_{5000} , where the density contrast, $\rho(r)/\rho_c$, is 2500 and 5000, respectively (see Appendix). R_{image} is the maximum radius analysed.

et al. adopt the NFW profile (Navarro et al. 1997) for the total mass.

Finally we must stress that although the new estimates of the X-ray dynamical mass presented here are higher than those given in Paper I, they are still a factor, ~ 5 – 10 , below the virial mass estimates (VME) obtained by the application of the Virial Theorem to the galaxies. As discussed in Paper I, in the case of equilibrium, the dynamical mass by the VME should not be overestimated by more than 20%. We are thus lead to interpret the large discrepancies between the virial theorem and the X-ray dynamical mass estimates as evidence of a non-equilibrium state of cluster. Indeed, in Paper I we had already remarked the incidence of substructures in the galaxy projected distribution and of a large scale gradient of their line-of-sight velocities, which are strong signs that the cluster is out of equilibrium. We will discuss more in the following Section.

5. Cluster merging

Figure 7 shows the temperature and metallicity measured in four quadrants: NW, NE, SW, and SE region.

As it can be seen from this figure, there is a hint of a temperature gradient in the North–South direction. Also, there is a possible metallicity gradient in the NE–SW direction. However, within their error bars, those values are also compatible with an isotropic distribution. In Paper I, it has been shown that there is a large-scale gradient of the local mean velocities of the cluster galaxies. The general direction of this velocity

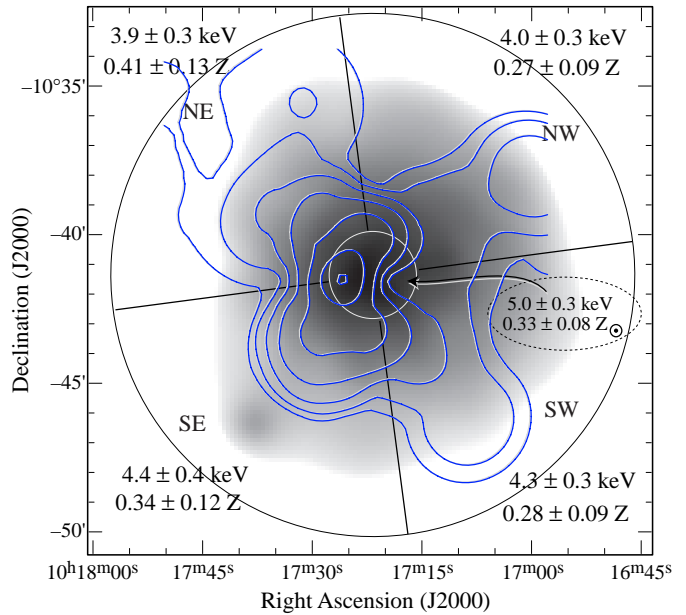


Fig. 7. Temperature and metallicity in the regions central, NW, NE, SW, and SE, superposed on the combined MECS 2 and 3 adaptively smoothed image. The solid lines mark the cluster regions for which we carried out the spectroscopic analysis. The central $1'28''$ region corresponds to the innermost point in Figs. 3–5. The isocontours (spaced logarithmically) of the projected density map of galaxies kinematically bound to the cluster and with $B < 19$, are superposed to the X-ray image (cf. Fig. 1 in Paper I).

gradient is approximately the same as the temperature gradient hinted in Fig. 7.

Moreover, we observe that the irregular galaxy distribution does not coincide with the X-ray emission map. There is also a discrepancy between the dynamical mass determined by the virial theorem and X-ray gas emission (see Sect. 4). These features are recognized as signs of past mergers between subclusters. We may infer that, in this case, the merger is not so recent, since the X-ray distribution already shows a fairly axial symmetric distribution.

From another perspective, as discussed in Paper I and shown in Fig. 7, the galaxy distribution presents a substructure located at a projected distance of $760 h_{50}^{-1}$ kpc (~ 8 arcmin) NW of the centre, with a velocity dispersion of 380_{-70}^{+120} km s $^{-1}$ inside a circular radius of $285 h_{50}^{-1}$ kpc (~ 3 arcmin). This was interpreted as a group falling into the core of Abell 970. Taking into account the mean velocities of the central region, 17600 ± 118 km s $^{-1}$, and NW substructure, 17834 ± 135 , (cf. Paper I), then we can roughly assume that the substructure may be falling with a velocity in the range 50 – 350 km s $^{-1}$. Consequently, with such a crude estimate, we conclude that a merger will occur in the next 2–6 Gyr.

Another possibility is that this substructure has already passed near the cluster centre, disturbing the overall cluster equilibrium, and will come back again to the cluster central regions, before being destroyed.

Therefore, the galaxy distribution (spatial and velocity) together with the X-ray isophote off-centre and possible temperature gradient suggest that Abell 970 is between two merger

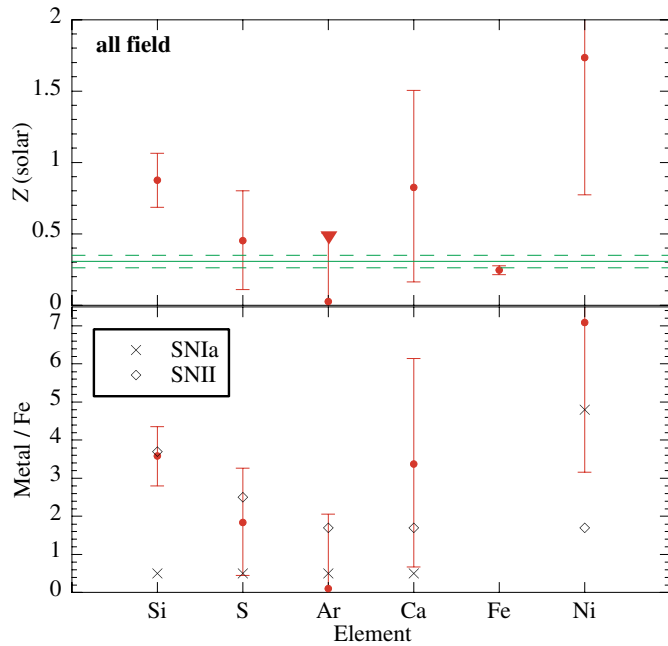


Fig. 8. Top: Individual metallicities of Fe, Ni and α -elements obtained with an absorbed vMEKAL model. The solid horizontal line is the mean metallicity obtained with an absorbed MEKAL model; the dashed lines correspond to 1σ confidence level, except for argon which is an upper limit. Bottom: abundance ratios between metal and iron. Also shown are the predicted approximate yields of Type II SNe (diamonds) and Type Ia SNe (crosses), based on Dupke & White (2000).

events: the first one that is responsible for the offset between the X-ray and galaxy distributions; the second one will occur when the NW substructure falls into the central part of the cluster.

6. Metal abundances

The pattern observed in the heavy-element abundances of the intracluster medium provides important clues on the types of galaxies and stars producing these elements and the mechanisms responsible for their removal from the interstellar medium of their parent galaxies.

Using the whole 8 arcmin field ($0.33\text{--}0.4 r_{200}$, see Sect. 4.2), we fitted a plasma model with the following metal abundances as free parameters: Si, S, Ar, Ca, Fe, and Ni. The following element abundances were fixed to $0.3 Z_{\odot}$: He, C, N, O, Ne, Na, Mg, and Al.

The result is summarized in Fig. 8, which shows the over-abundance of α -elements relative to iron. Both the silicon and sulphur abundances are consistent with the typical values found recently by Finoguenov et al. (2001b) and are slightly above the value determined by Fukazawa et al. (1998) for a cluster of ~ 4.5 keV (which is $\sim 0.6\text{--}0.8 Z_{\odot}$, cf. Fig. 2 of their paper). These authors have shown that the Si abundance correlates with the gas temperature, probably due to a decreasing contribution from Type II supernovae (SNe) in poorer clusters, where the metals produced by this type of SN could escape from the potential well.

The over-density of α -elements displayed in Fig. 8 is characteristic of the Type II SNe yield. This suggests that the intracluster gas enrichment was caused by ejected gas that was synthesized in massive stars, the progenitors of Type II SNe. Present day elliptical galaxies show a very low rate of Type II SNe, implying that the enrichment of the intracluster gas probably occurred early in the history of the cluster.

However, we cannot discard other possible enrichment mechanisms. In fact, as shown in Fig. 8, the abundance ratio Ni/Fe is higher than expected for Type II SNe, but it is in agreement with the expected Type Ia SNe yield. This finding supports models where part of the metals in the ICM were produced by Type Ia SNe and which were either blown from the galaxies by SNe explosions or swept by ram-pressure (e.g., Dupke & White 2000). However, the Ni abundance is uncertain. Moreover, the abundance ratio Si/Fe (with its smaller error bars) is incompatible with enrichment by Type Ia SNe.

7. Conclusions

We have presented here a new X-ray observation of Abell 970. Below we summarize our main results.

- Although Abell 970 presents rather symmetric X-ray isophotes (except for an excess emission that is probably due to an active galactic nucleus, triggered by a galaxy merger), the peak of the X-ray distribution does not coincide with the centre of the galaxy distribution.
- The temperature profile presents a strong radial gradient at ~ 6 arcmin, which is barely compatible with adiabatic stable equilibrium. The metal abundances are approximately constant up to 8 arcmin. There are indications of non radial temperature and metallicity gradients (but with confidence levels less than 1σ), notably in the NW–SE direction, the same direction where a gradient in velocity is observed (cf. Paper I).
- With the two findings above, we arrive at two possible scenarios for the recent history of Abell 970: either it has suffered a recent merger with a subcluster, or the NW substructure (detected only from galaxy counts) has recently passed through the centre of Abell 970. In both cases, the dynamical stability was disrupted, but the cluster is now approaching virial equilibrium again. The observed galaxy distribution suggests that this equilibrium will be broken in a few Gigayears when the nearby NW substructure will fall (again?) into the core of Abell 970.
- We have derived an abundance ratio between α -elements/iron greater than 1 and the iron metallicity is lower than the mean metallicity. The over abundance of α -elements may be an indication of an early enrichment of the ICM by Type II supernovae in elliptical galaxies. Part or the metals in the ICM could be produced by Type Ia SNe, but the evidence from the Ni/Fe ratio is not strong.
- We have found a central X-ray absorption excess compared to the absorption expected from the mean galactic value of N_{H} , although the statistical error bars are large. The central N_{H} excess could be associated with the central galaxies

and/or with cooled intra-cluster material (either in molecular form and/or dust grains). The latter may be preferable since the core of the cluster is dominated by early-type galaxies. However, this result probably depends on the spectral modelling (see, e.g., McCarthy et al. 2002).

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Appendix A: Derivation of r_δ

In Paper I, we give a formula for r_δ , which is defined by:

$$\delta \equiv \frac{\bar{\rho}(r_\delta)}{\rho_c(z)}, \quad (\text{A.1})$$

where $\bar{\rho}(r_\delta) \equiv M_{\text{dyn}}(r_\delta)/(4\pi r_\delta^3/3)$ is the mean density of the cluster inside r_δ and $\rho_c(z)$ is the critical density of the Universe at the redshift of the cluster.

We take the dynamical mass corresponding to a β -model ICM and polytropic temperature profile, Eq. (6), and the critical density given by:

$$\rho_c(z) = \frac{3H^2(z)}{8\pi G}; \quad H(z) = H_0 f(z, \Omega_M, \Omega_\Lambda), \quad (\text{A.2})$$

with

$$f^2(z, \Omega_M, \Omega_\Lambda) = \frac{\Omega_\Lambda + \Omega_M(1+z)^3 - (\Omega_\Lambda + \Omega_M - 1)(1+z)^2}{\Omega_\Lambda + \Omega_M(1+z)^3 - (\Omega_\Lambda + \Omega_M - 1)(1+z)^2}. \quad (\text{A.3})$$

Combining these equations and solving for r_δ we obtain:

$$\left(\frac{r_\delta}{r_c}\right)^2 + 1 = \left[\frac{2.3 \times 10^8 \beta \gamma T_0}{\delta h_{50}^2 f^2(z, \Omega_M, \Omega_\Lambda) \mu r_c^2} \right]^{1/[1+3\beta(\gamma-1)/2]}, \quad (\text{A.4})$$

where the cluster central temperature, T_0 , is given in keV and the core radius, r_c , in kpc. For $\delta \lesssim 500$, $r_\delta \gg r_c$, and Eq. (A.4) is simplified:

$$\frac{r_\delta}{r_c} \simeq \left[\frac{2.3 \times 10^8 \beta \gamma T_0}{\delta h_{50}^2 f^2(z, \Omega_M, \Omega_\Lambda) \mu r_c^2} \right]^{1/[2+3\beta(\gamma-1)]}. \quad (\text{A.5})$$

When $\gamma = 1$, the isothermal case is obtained and T_0 is the mean temperature.

For $z \ll 0.1$ we can use the approximation:

$$f^2(z, \Omega_M, \Omega_\Lambda) \simeq 1 + (2 - 2\Omega_\Lambda + \Omega_M)z. \quad (\text{A.6})$$

References

- Allen, S. W. 2000, MNRAS, 315, 269
 Allen, S. W., Schmidt, R. W., & Fabian, A. C. 2002, MNRAS, 334, L11
 Balucinska-Church, M., & McCammon, D. 1992, ApJ, 400, 699
 Boella, G., Butler, R. C., Perola, G. C., et al. 1997a, A&AS, 122, 299
 Boella, G., Chiappetti, L., Conti, G., et al. 1997b, A&AS, 122, 327
 Cavaliere, A., & Fusco-Femiano, R. 1976, A&A, 49, 137
 Cavaliere, A., Menci, N., & Tozzi, P. 1999, MNRAS, 308, 599
 Cusumano, G., & Mineo, T. 1998, On-flight calibration of the MECS vignetting
 De Grandi, S., & Molendi, S. 1999, ApJ, 527, L25
 De Grandi, S., & Molendi, S. 2002, ApJ, 567, 163
 Dickey, J. M., & Lockman, F. J. 1990, ARA&A, 28, 215
 Dos Santos, S., & Doré, O. 2002, A&A, 383, 450
 Dupke, R. A., & White, R. E., III, 2000, ApJ, 537, 123
 Einasto, M., Tago, E., Jaaniste, J., Einasto, J., & Andernach, H. 1997, A&AS, 123, 119
 Evrard, A. E., MacFarland, T. J., Couchman, H. M. P., et al. 2002, ApJ, 573
 Finoguenov, A., Reiprich, T. H., & Böhringer, H. 2001a, A&A, 368, 749
 Finoguenov, A., Arnaud, M., & David, L. P. 2001b, ApJ, 555, 191
 Fiore, F., Guainazzi, M., & Grandi, P. 1999, Cookbook for BeppoSAX NFI spectral analysis, version 1.2, www.sdc.asi.it/pub/sax/doc/software_docs/saxabc_v1.2.ps.gz
 Fukazawa, Y., Makishima, K., Tamura, T., et al. 1998, PASJ, 50, 187
 Irwin, J. A., Bregman, J. N., & Evrard, A. E. 1999, ApJ, 519, 518
 Irwin, J. A., & Bregman, J. N. 2000, ApJ, 538, 543
 Irwin, J. A., & Bregman, J. N. 2001, ApJ, 546, 150
 Jones, C., & Forman, W. 1999, ApJ, 51, 65
 Kaastra, J. S., & Mewe, R. 1993, A&AS, 97, 443
 Kalinkov, M., Valtchanov, I., & Kuneva, I. 1998, A&A, 331, 838
 Komatsu, E., & Seljak, U. 2001, MNRAS, 327, 1353
 Lacey, C., & Cole, S. 1993, MNRAS, 262, 627
 Lea, S. M., Silk, J., Kellogg, E., & Murray, S. 1973, ApJ, 184, 105L
 Liedahl, D. A., Osterheld, A. L., & Goldstein, W. H. 1995, ApJ, 438, L115
 Lima Neto, G. B., Pislár, V., & Bagchi, J. 2001, A&A, 368, 440
 Loken, C., Melott, A. L., & Miller, C. J. 1999, ApJ, 520, L5
 Markevitch, M., Vikhlinin, A., Forman, W. R., & Sarazin, C. L. 1999, ApJ, 527, 545
 McCarthy, I. G., West, M. J., & Welch, G. A. 2002, ApJ, 567, 762
 Navarro, J. F., Frenk, C. S., & White, S. D. M. 1997, ApJ, 490, 493
 Parmar, A. N., Martin, D. D. E., Bavdaz, M., et al. 1997, A&AS, 122, 309
 Parmar, A. N., Oosterbroek, T., Orr, A., et al. 1999, A&AS, 136, 407
 Pislár, V., Durret, F., Gerbal, D., Lima Neto, G. B., & Slezak, E. 1997, A&A, 322, 53
 Schwöpe, A., Hasinger, G., Lehmann, I., et al. 2000, AN, 321, 1
 Sodré, L., Jr., Proust, D., Capelato, H. V., et al. 2001, A&A 377, 428 (Paper I)
 Suto, Y., Sasaki, S., & Makino, N. 1998, ApJ, 509, 544
 White, D. A., Jones, C., & Forman, W. 1997, MNRAS, 292, 419
 White, S. D. M., & Rees, M. J. 1978, MNRAS, 183, 341
 White, D. A., & Fabian, A. C. 1995, MNRAS, 273, 72
 Xue, Y. J., & Wu, X. P. 2000, ApJ, 538, 65